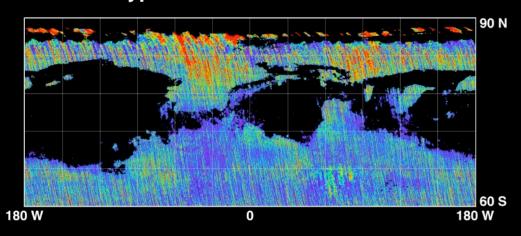
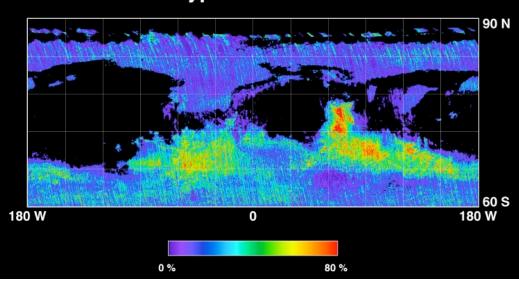


### In the beginning there was Type I and Type II ....

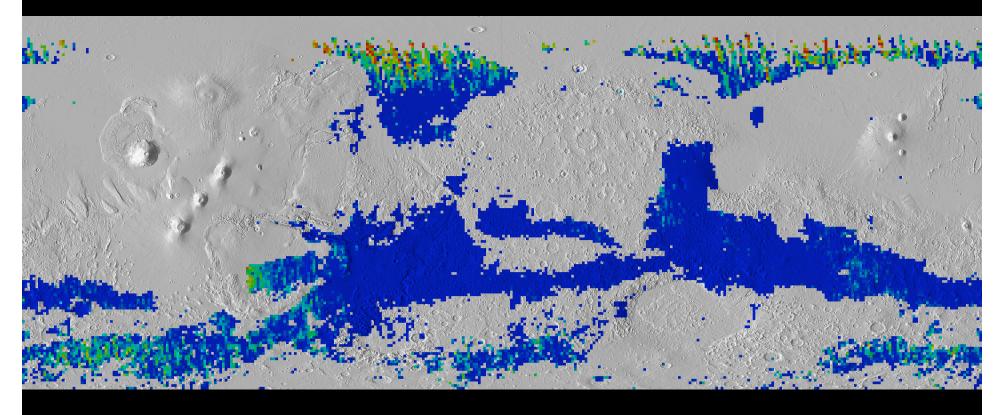
Type 2 – Basaltic Andesite



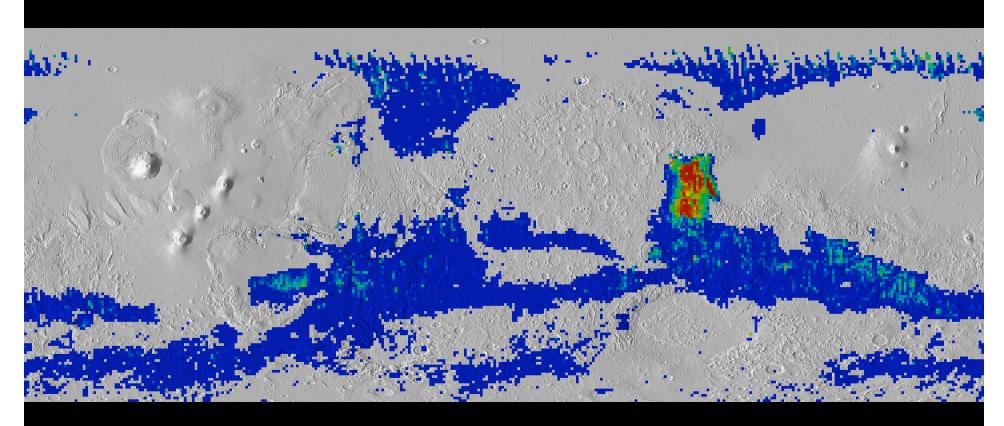
Type 1 - Basalt



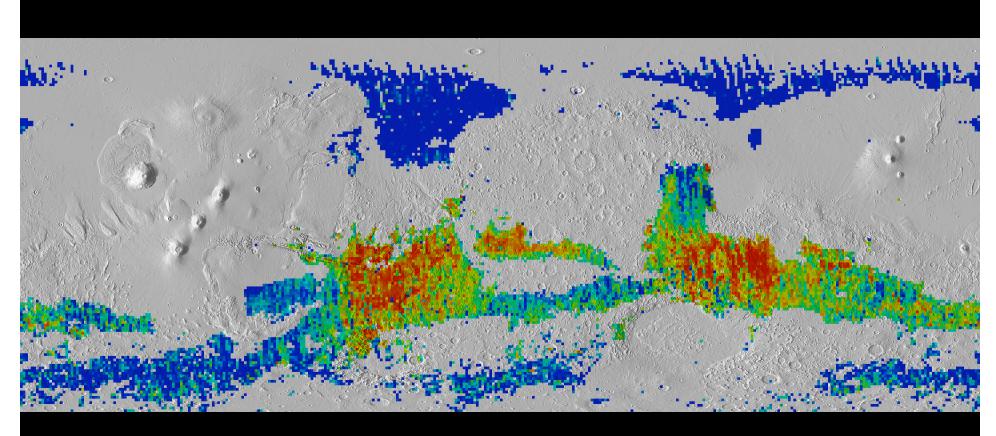
# Today ...



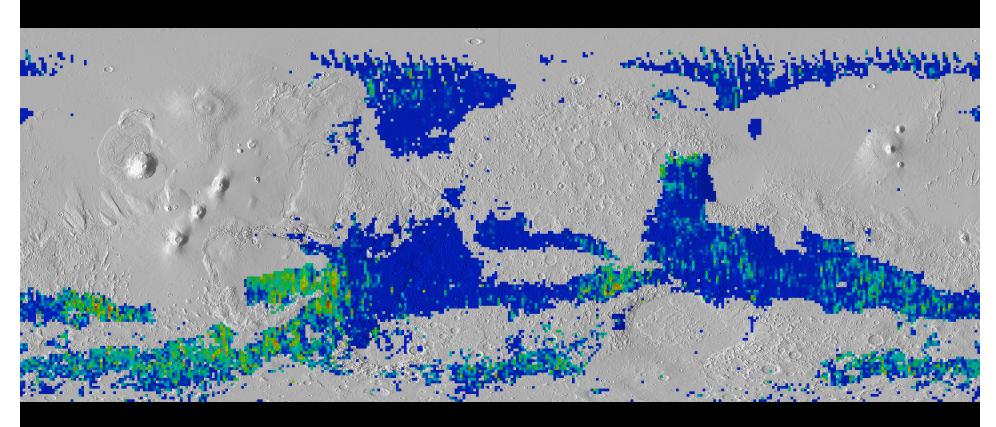
Group 1



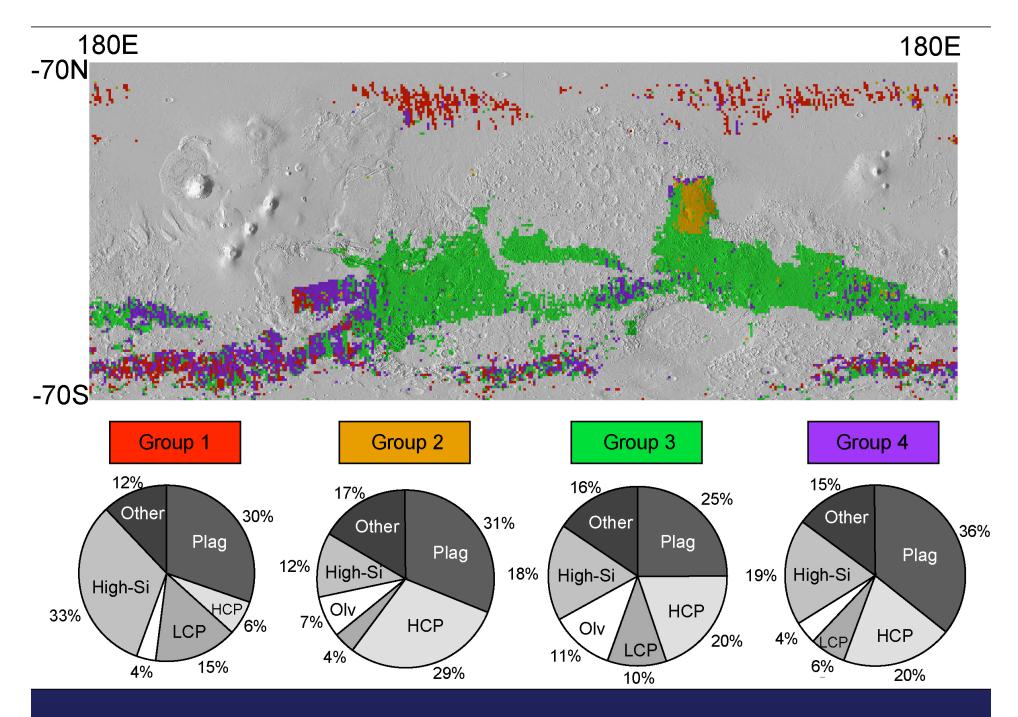
Group 2



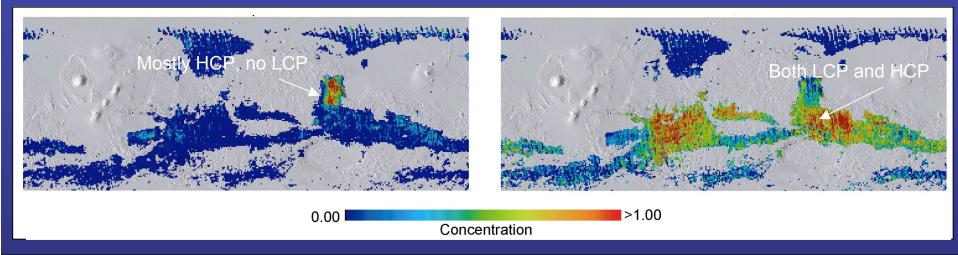
Group 3

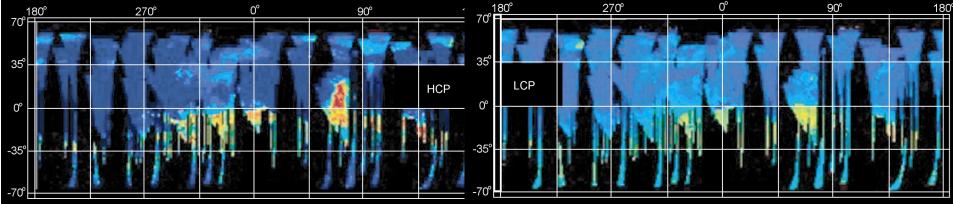


Group 4



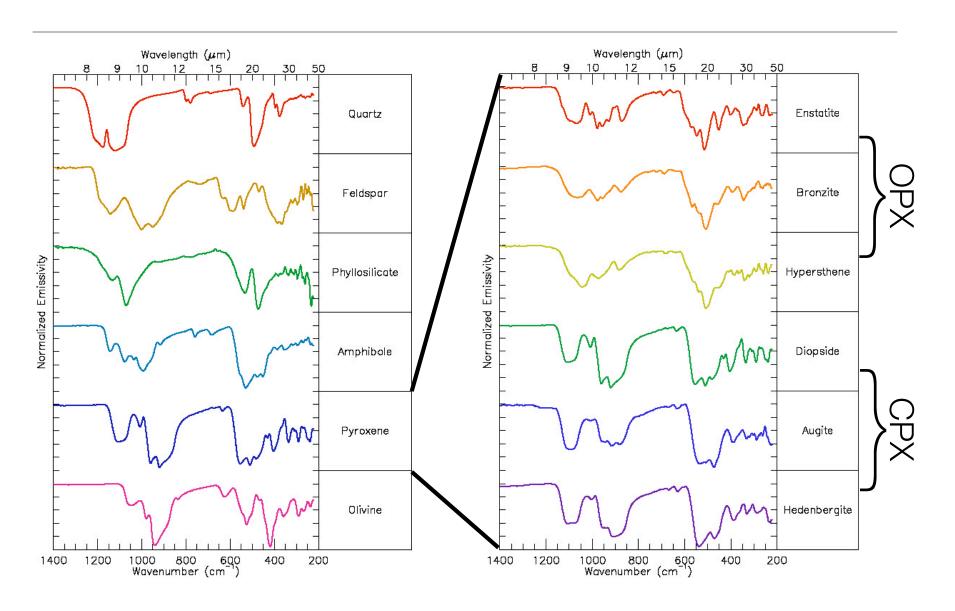
# Pyroxene compositions: TES-OMEGA comparison



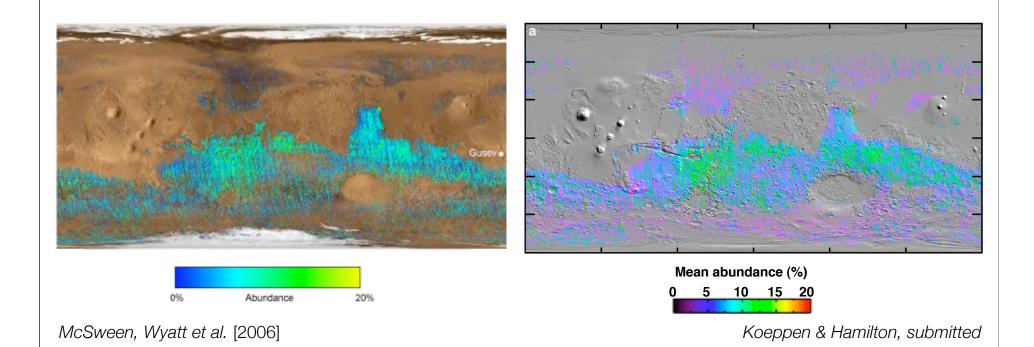


Bibring et al., Science, [2005]

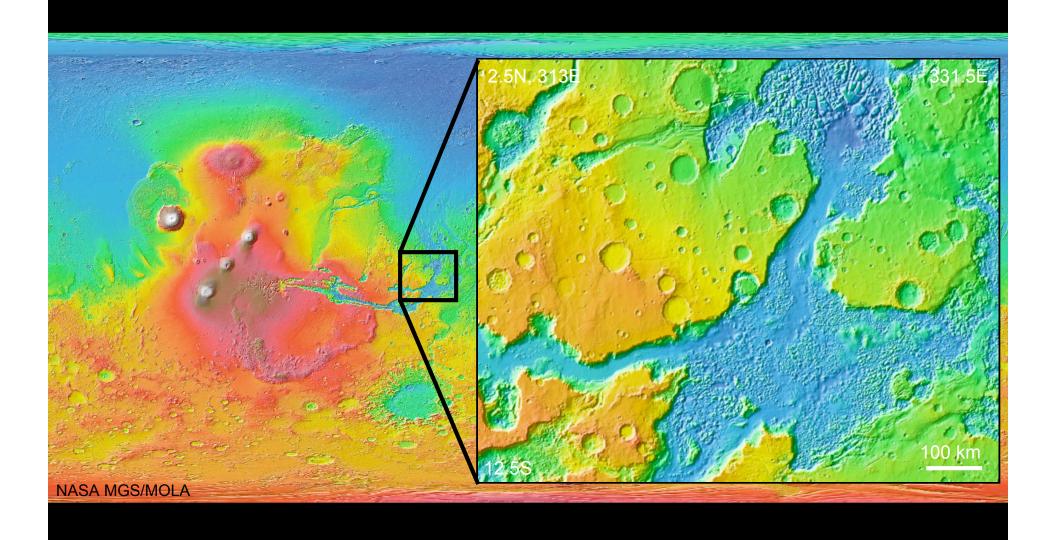
#### Improved libraries of pyroxenes and olivines in the infrared

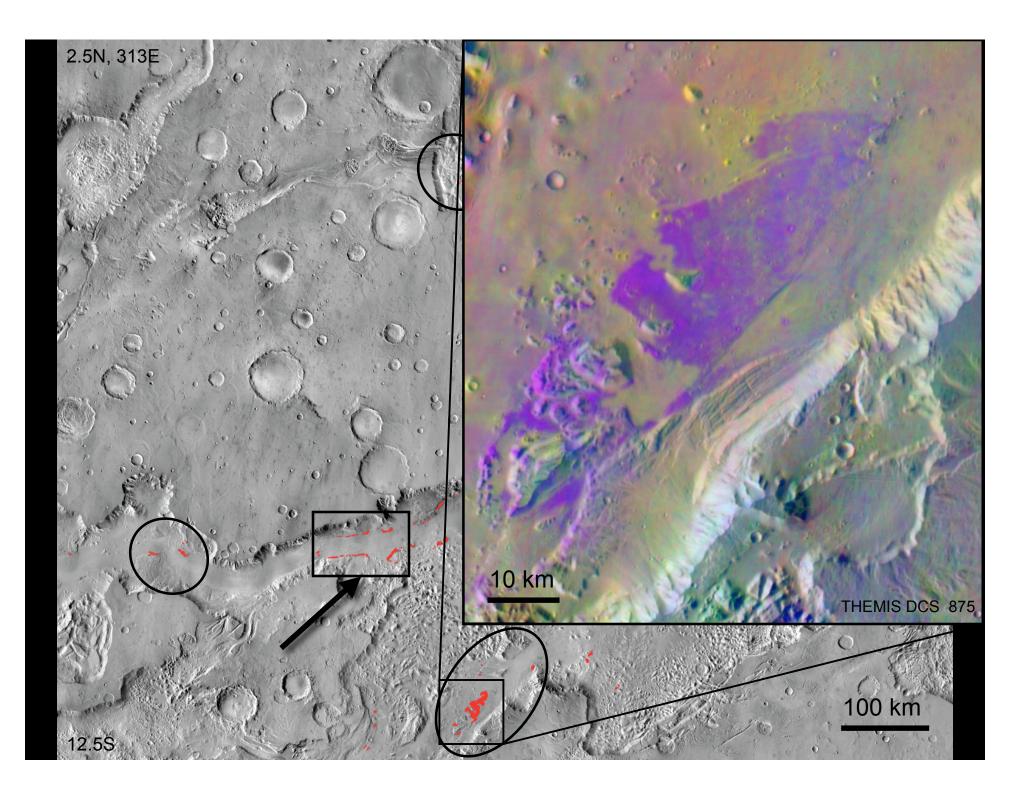


#### TES-derived surface abundance of olivine

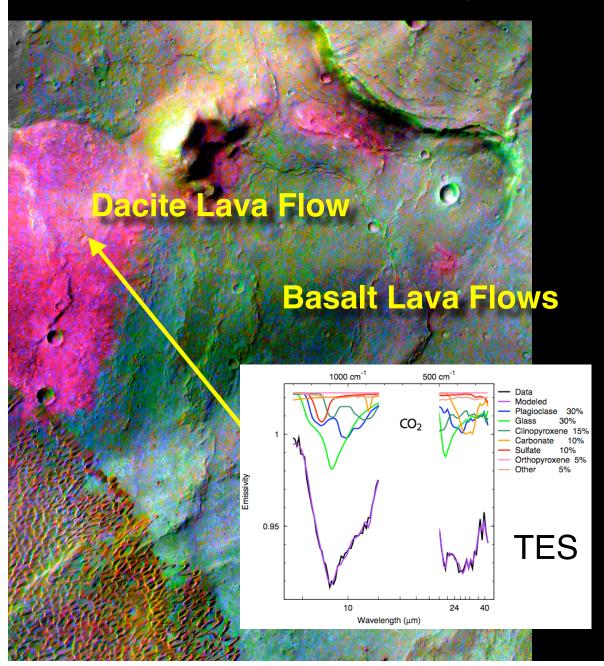


# Olivine Basalt in Eastern Valles Marineris

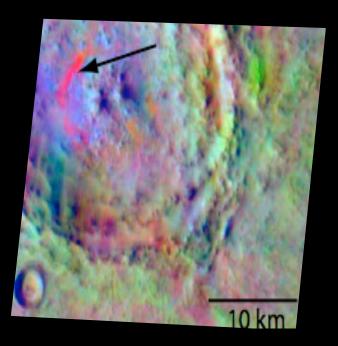


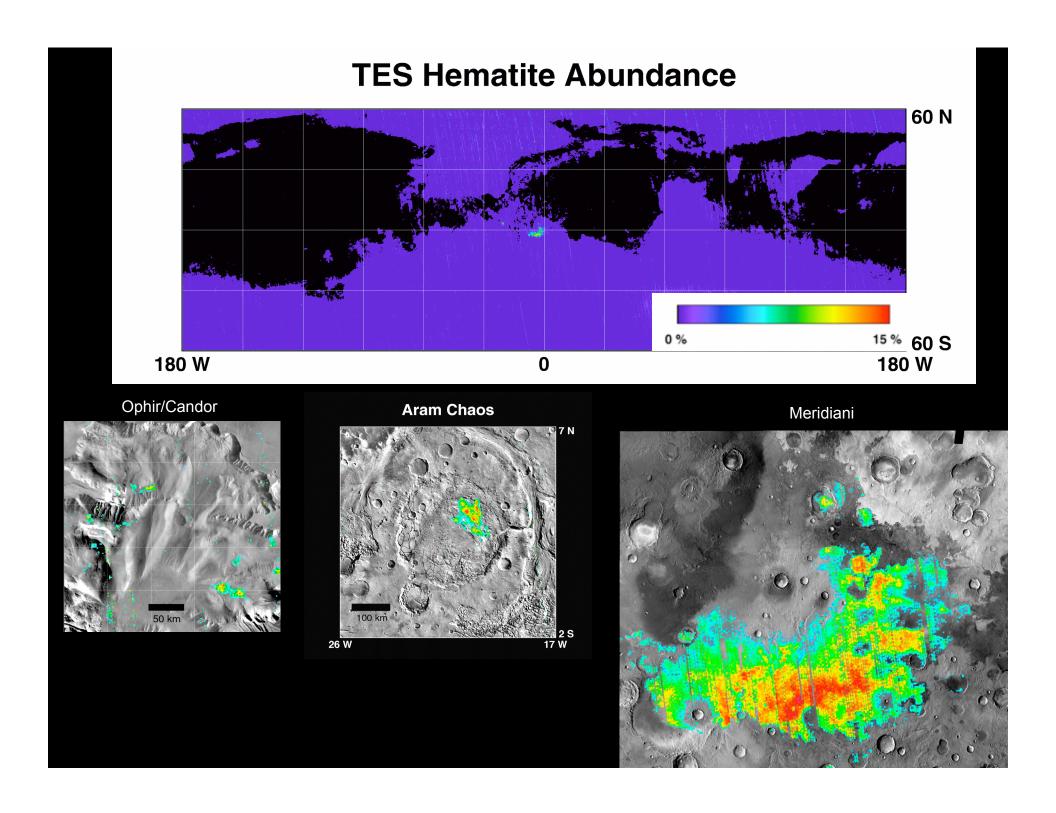


#### THEMIS Multi-spectral IR



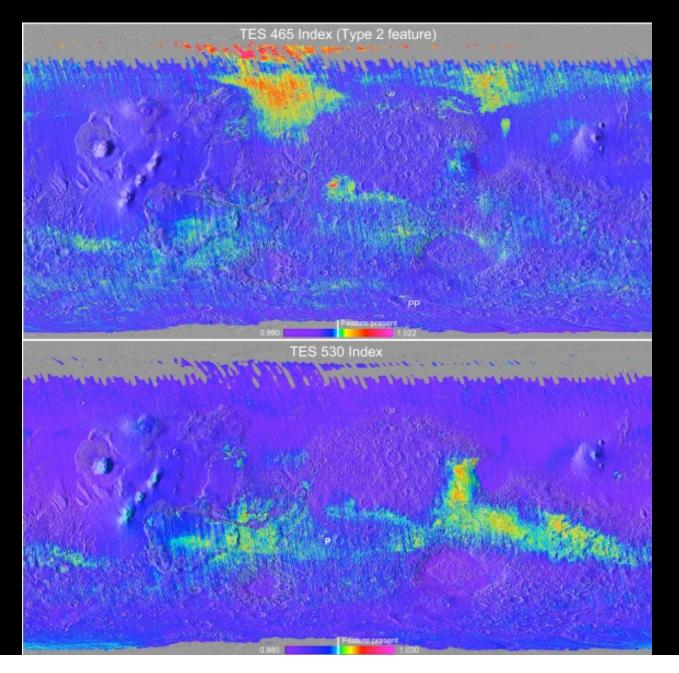
#### **Granitoid Rocks**





#### Carbonates 1500 cm<sup>-1</sup> 1000 cm<sup>-1</sup> 500 cm<sup>-1</sup> 0.95 Emissivity 0.9 Carbonate Mini-TES **Bound Water** 0.85 **Orbiting MGS/TES** 10 6 24 Wavelength (µm)

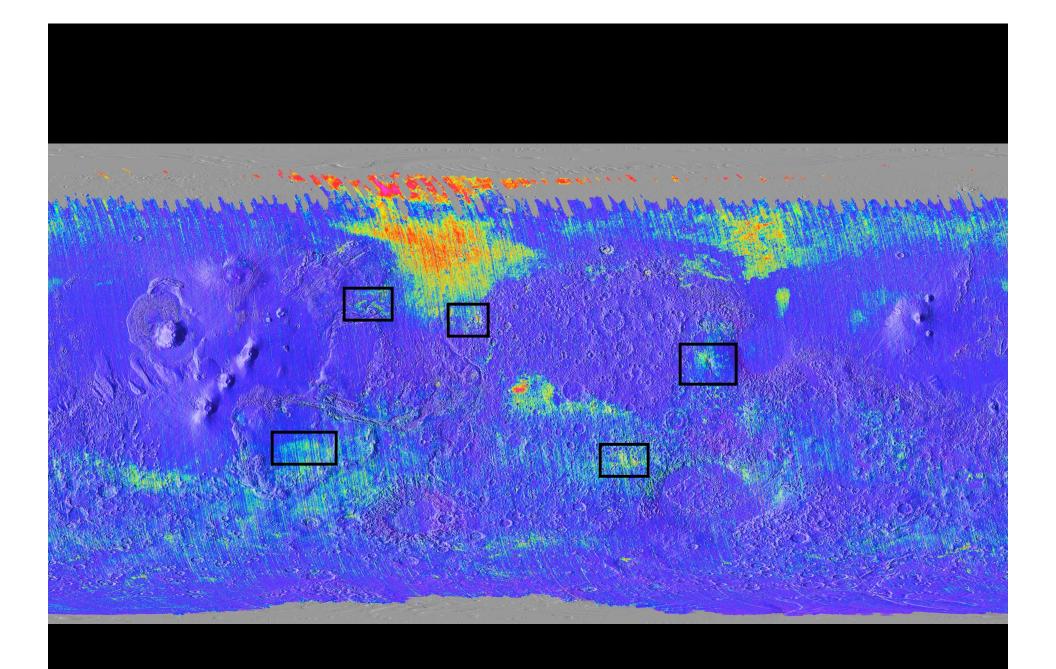
## Phyllosilicates

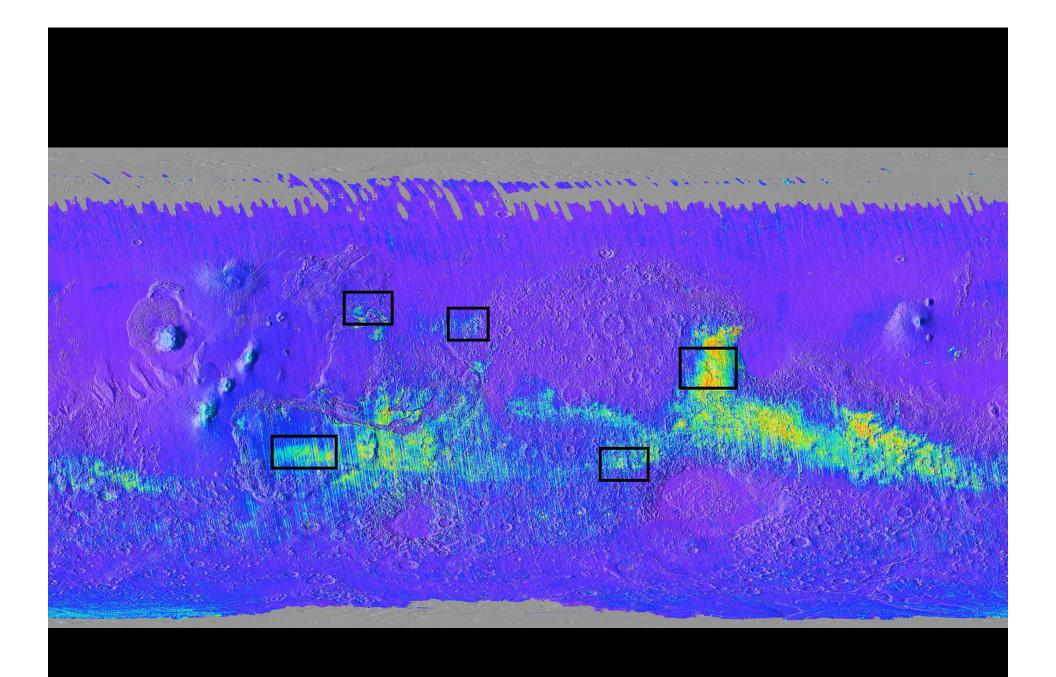


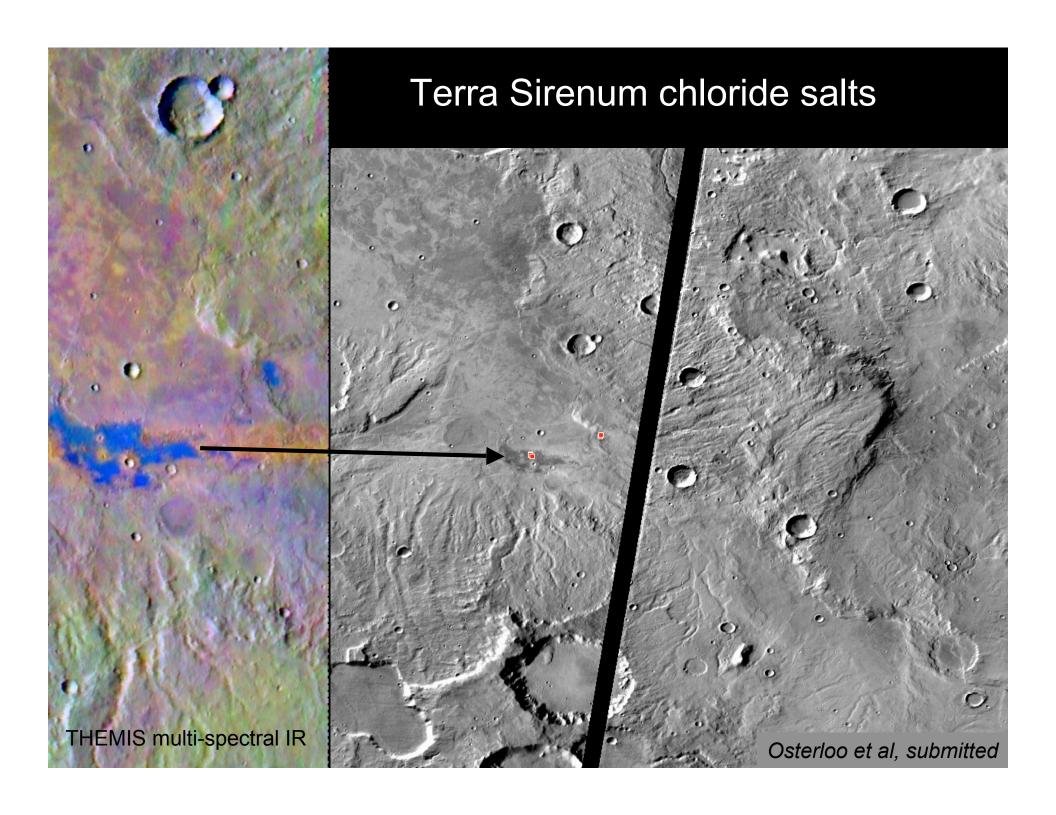
Steve Ruff's 465 cm<sup>-1</sup> index

Both present = crystalline clay

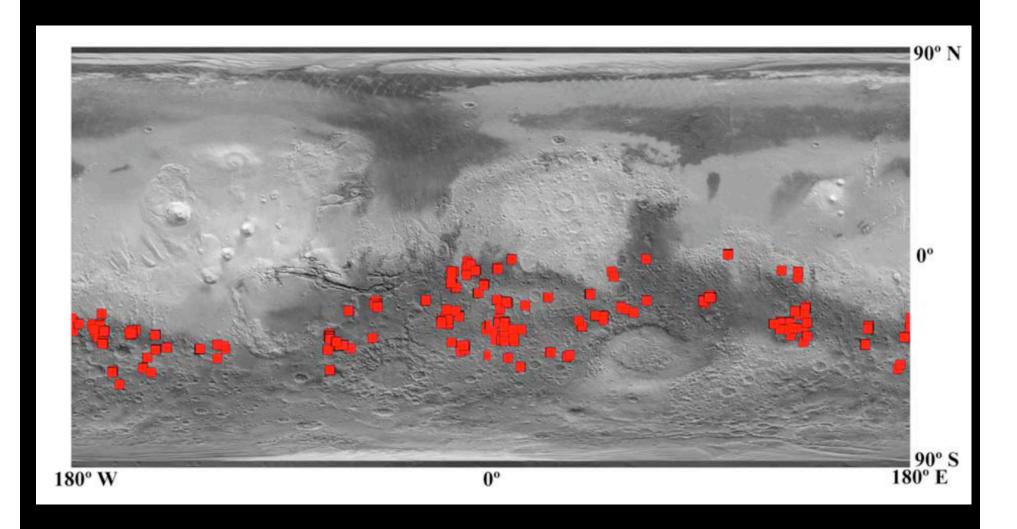
530 cm<sup>-1</sup> index

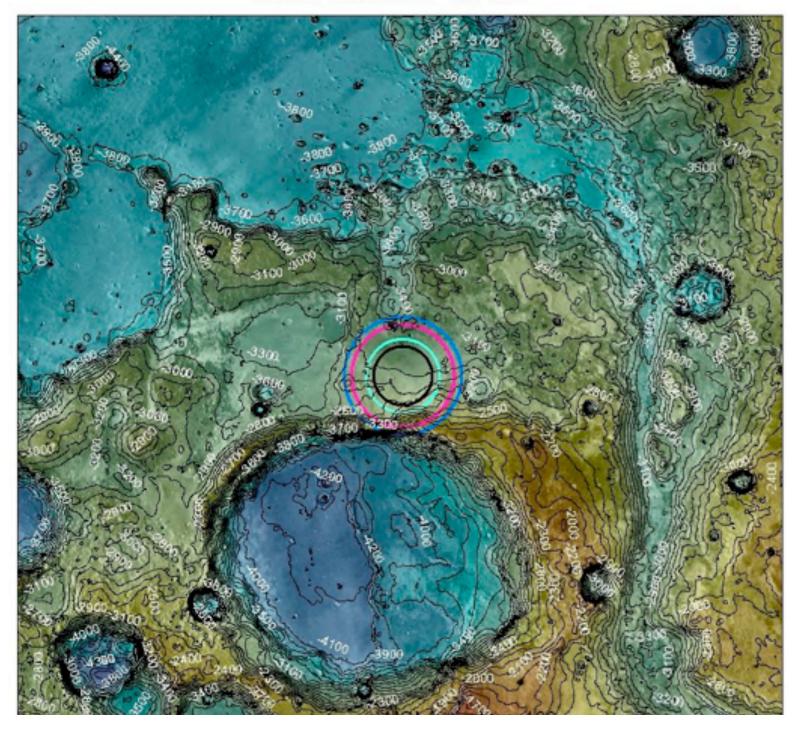


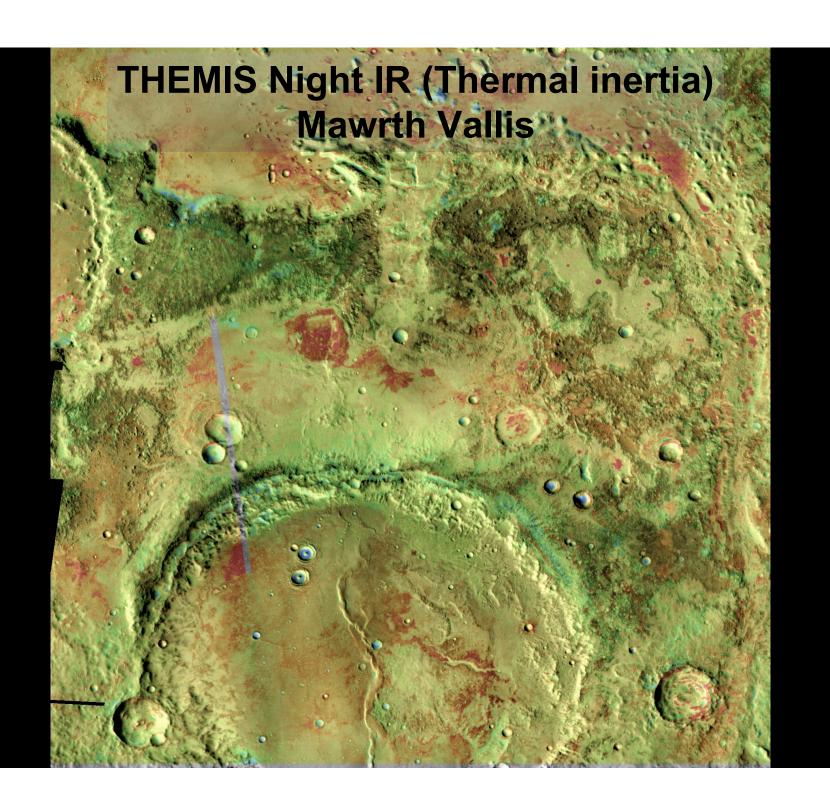




## Salt Sites







## Summary

- Knowledge of martian mineralogy from TES and THEMIS data has improved substantially over the past 5 years
- Mars is a volcanic planet with diverse compositions
- Weathering has occurred, but extensive exposures of unweathered rock and regolith are present
  - Water-related environments need to be kept in this context
- Clays (and sulfates) appear from TES and THEMIS data to be present at abundances of <~10%</li>
- Carbonates are present but in dust
- Chloride salts appear to have been detected
- Bedrock is present Mars is still actively (physically) eroding